

1 Title: METHOD AND APPARATUS FOR IMAGING A FIELD OF REGARD BY  
2 SCANNING THE FIELD OF VIEW OF AN IMAGING ELECTRO-  
3 OPTICAL SYSTEM IN A SERIES OF CONICAL ARCS TO COMPENSATE  
4 FOR IMAGE ROTATION

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7 **STATEMENT CONCERNING GOVERNMENT INTEREST**

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10 provided for by the terms of contract no. NAS5-32650 awarded by the National  
11 Aeronautics and Space Administration (Goddard Space Flight Center).

12 **BACKGROUND OF THE INVENTION**

13 **1. Field of the Invention**

14 The present invention relates generally to the field of imaging a two-dimensional  
15 field of regard. More particularly, the present invention relates to imaging of the full-  
16 Earth disk by a spacecraft that scans the field of view of an imager across the full Earth  
17 disk.

18 **2. Background Information**

19 One of the most common uses for artificial planetary satellites is to produce images  
20 of the planet's surface. Many Earth-orbiting satellites capture images of Earth for  
21 purposes ranging from military intelligence to weather forecasting. Orbital imaging for  
22 weather forecasting and for scientific purposes demands images of vast areas of the Earth  
23 at once.

24 It is common to image a planet (e.g., Earth) from space using an imaging electro-  
25 optical sensor constructed from a telescope that collects radiation from a remote source

1 and brings it to focus on one or more focal plane arrays (FPA's) with each FPA containing  
2 many detector elements. A scanning sensor moves the image of the scene over one or  
3 more FPA's, each FPA usually having many detector elements perpendicular to the  
4 direction of the scan motion. Each element converts the radiation from an instantaneous  
5 field of view (IFOV) in the scene into an electronic signal. The image of the scene in the  
6 spectral band of the sensor is reconstructed from these electronic signals.

7 The angular field of view (FOV) of the telescope multiplied by its effective focal  
8 length (EFL) equals the dimensions of the telescope's focal plane, which contains the  
9 FPA. For application in which an imaging sensor must cover a two-dimensional field of  
10 regard (FOR) that exceeds the telescope's FOV, a plane scan mirror may be located in  
11 front of the imaging sensor's telescope to scan the FOV across the FOR. For example, the  
12 Earth subtends a circle approximately  $17.4^\circ$  in diameter from geosynchronous altitude. An  
13 instrument that is capable of imaging the full-Earth disk must have a field of regard (FOR)  
14 that not only includes this full-Earth disk, but also allows it to view deep space to measure  
15 the background signal in each channel. Most multispectral instruments that image the  
16 Earth from this altitude use a large, reflective telescope with a field-of-view (FOV) that is  
17 much smaller than the required FOR. A two-dimensional raster scanning procedure is  
18 required to cover the FOR, and is usually implemented with a plane mirror in front of the  
19 telescope's aperture.

20 A number of geosynchronous weather satellites, including the EUMETSAT and  
21 GOES-1 through GOES-7 satellites, are "spinners" that rotate about the north-south axis.  
22 The imager on each of these spinning satellites has a telescope that is aligned with the  
23 north/south axis of the spacecraft. A plane mirror with a single rotational axis,  
24 perpendicular to the spin axis, reflects the optical axis of the telescope towards the Earth.

1 The spacecraft's rotation scans the line of sight (LOS) in the east/west direction. To form  
2 a two-dimensional map of the Earth, the plane mirror is only required to step in the  
3 north/south direction. The main disadvantage of a spinning satellite is that it only allows  
4 the imager to view the Earth's surface for less than 5% of its total duty cycle.

5 Beginning with GOES 8, the geostationary weather satellites operated by the  
6 United States (developed for NOAA by LORAL with instruments from ITT) have been  
7 three-axis stabilized. In this configuration, the imager continuously points toward the  
8 Earth, permitting it to operate at a high duty cycle and to be far more flexible than a  
9 spinning satellite in imaging arbitrary areas of the Earth's surface. These prior art GOES  
10 imagers routinely produce 3000km by 5000km images of the contiguous United States  
11 (CONUS) and 1000km by 1000km images of severe storms. Scanning is performed by a  
12 plane scan mirror mounted on a two-axis gimbal. Rotation of this mirror about the inner  
13 gimbal axis scans the LOS in the east/west direction. Between scan lines, incremental  
14 rotations about the outer gimbal axis move the LOS from north to south. When scanning  
15 the Equator, the GOES imager projects its detector arrays onto the Earth's surface in the  
16 optimal manner, with the cross-scan axis of each detector array (its long axis) projected in  
17 the north-to-south direction. When scanning north of the Equator, the projection of this  
18 axis is tilted in the northeast- to-southwest direction; when scanning south of the Equator,  
19 the projection is tilted northwest-to-southeast. The tilt angle varies from zero at the  
20 Equator to  $8.7^\circ$  at the North and South Poles. This phenomenon, known as image rotation,  
21 is an intrinsic problem in a two-axis scanning system that uses a single scan mirror in  
22 object space. See J. J. Shea, "Image correction via lunar limb knife-edge OTF's", *Proc.*  
23 *SPIE, Earth Observing Systems III*, vol. 3439, 19-21 July 1998, pp 165-186.

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Referring to Fig. 1, the geometric configuration of the GOES 8 & 9 imager is illustrated. The GOES convention for spacecraft coordinates is portrayed by a set of orthogonal coordinates **10** wherein  $+x$  = east,  $+y$  = south, and  $+z$  = nadir. For simplicity of illustration, the GOES telescope is represented as a single lens **12** and crossed axes **14** represent the image as presented at the focal plane array. Note that the lens **12** inverts the image of the axes **14**. The telescope's optical axis **16** points due east along the x-axis. The scan mirror **18** is an optical flat with an elliptical cross section and has a reflective surface (not visible from the viewpoint of Fig. 1) that directs light from the Earth (shown in phantom) into the telescope **12**. The scan mirror **18** is mounted on a first axle **20** that provides an inner axis of rotation with respect to which the inner gimbal angle (iga) is measured. The inner axis of rotation is coincident with the short dimension of the ellipse and perpendicular to the normal vector of the mirrored surface. The first axle **20** permits the scan mirror **18** to rotate about the inner axis of rotation with respect to a yoke **22**. The yoke **22** has a second axle **24** that is perpendicular to the first axle **20**. The second axle **24** lies along the extension of the optical axis **16** of the telescope **12**, along the x-axis, and allows the yoke **22** to rotate about this outer axis of rotation that is fixed with respect to the telescope **12**, and with respect to which the outer gimbal angle (oga) is measured. The orientation of the first axle **20** always remains perpendicular to the x-axis, but rotates in the y-z plane when the yoke **22** is pivoted about the second axle **24**.

Referring to Fig. 2, projections **42, 44, 46, 52, 54, 56** of the crossed axes in the focal plane **14** onto the Earth's surface **30** are illustrated. The line with the arrowhead **14'** is parallel to the z-axis and corresponds to the along-scan direction of the array. The line with the circle **14''** is parallel to the y-axis and corresponds to the cross-scan direction of the array.

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1 The intersection of the crossed axes is projected onto the equator 32 when the  
2 position of the yoke 22 on the outer axle 24 aligns the inner axle parallel to the y-axis.  
3 This angle can be defined as the home position of the outer axle, at  $\text{oga} = 0$ . When the  
4 outer axle 22 is fixed at this position and the scan mirror 18 is rotated about its inner axle  
5 20, the projection 42, 44, 46 of the focal plane 14 is scanned along the equator 32. The y-  
6 axis of the focal plane remains perpendicular to the direction of the scan and the z-axis of  
7 the focal plane is projected along the direction of scan.

8 When the yoke 22 is rotated about the outer axle 24 in the  $+\text{oga}$  direction, the  
9 crossed axes in the focal plane 14 are projected 52, 54, 56 into the Northern Hemisphere.  
10 The array's projection 52, 54, 56 rotates clockwise, as viewed from space. When the oga  
11 remains fixed and the scan mirror 18 is rotated about the inner axle 20, the z-axis of the  
12 focal plane is no longer projected along the direction of scan but it is tilted away from this  
13 position by an angle equal to oga, reaching an angle of 8.7 at the North and South poles.  
14 When the  $\text{oga} < 0$  (not shown), the focal plane 14 is projected into the Southern  
15 Hemisphere with a counterclockwise rotation equal to the magnitude of the oga.

16 The GOES 8 and GOES 9 imagers scan the Earth in a series of alternating east-to-  
17 west and west-to-east lines proceeding from North to South. The GOES 8 and 9 imagers  
18 only maintain the optimal relationship between the direction of the opto-mechanical scan  
19 and the orientation of the array when they scan the equator. However, rotation of the  
20 mirror about a second axis rotates the projection of the focal plane onto the Earth's  
21 surface. This is a disadvantage of the prior art because it precludes use of a resolution  
22 enhancing technique known as Time Delay and Integration (TDI). The fundamental  
23 aspects of TDI technique are explained as follows, along with an explanation of why TDI  
24 is incompatible with the prior art systems that suffer from image rotation.

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1 TDI is a prior art technique in which a two-dimensional detector array is scanned  
2 like a linear array. TDI achieves the angular resolution of a single detector element in the  
3 array combined with a dwell time that is N times longer, where N is the number of  
4 columns in the along-scan direction. Referring to Fig. 3, TDI is illustrated, showing the  
5 projection of a 4x4 array onto object space, with the solid square, ■, representing a single  
6 charge packet. One each cycle, the charge packet in each detector element of the array  
7 accumulates photo-electric charges generated by absorption of radiation from the scene  
8 and is then transferred to the detector element in the same row that lies immediately to its  
9 right. The column on the far right of the array is also read out at the end of each cycle. At  
10 the end of cycle 4, the charge packet represented by the solid square contains the sum of  
11 the charges accumulated by the four detector elements in the second row: column: # 1 (the  
12 far left column) in cycle 1, column #2 in cycle 2, column # 3 in cycle 3, and column #4 in  
13 cycle 4.

14 The opto-mechanical scan motion of the LOS must be coordinated with the  
15 electronic scan of the charge packets so that the LOS from each charge packet corresponds  
16 to a single pixel in object space. The linear velocity of the electronic scan, from the  
17 leading edge to the trailing edge of the array, must match the linear velocity of the opto-  
18 mechanical scan. Referring to Fig. 4, the projection of the row or detector elements in  
19 object space is illustrated, where the right-to-left opto-mechanical scan cancels the left-to-  
20 right electronic scan, so that the image of a charge packet remains at a fixed angle in  
21 object space, corresponding to a fixed IFOV on the Earth's surface. The effective dwell  
22 time is equivalent to the dwell time of a single IFOV multiplied by N (N=4 in this  
23 example), so that radiometric signal-to-noise ratio (SNR) is enhanced by a factor of  $N^{1/2}$  (2

1 in this example) in comparison to a single column FPA scanned at the same opto-  
2 mechanical rate.

3 The goal of TDI is to accumulate all of the photo-electric charges in each charge  
4 packet from the same instantaneous field of view (IFOV) on the Earth's surface. To  
5 achieve this condition, the LOS of the present invention is scanned opto-mechanically so  
6 that the projection of the array moves horizontally, from left-to-right, with its velocity  
7 parallel to the rows of the array. At the same time, the array is electronically scanned so  
8 that a charge packet moves along each row from right-to-left. The electronic and opto-  
9 mechanical scan vectors are substantially equal in magnitude and opposite in direction to  
10 maintain image quality.

11 To achieve the desired match between opto-mechanical and electronic scans, two  
12 conditions need to be satisfied for all scan angles within the FOR. Ideally, the projection  
13 of the TDI axis of each FPA onto the Earth's surface must always lie along the direction of  
14 the opto-mechanical scan. If not, the pixels in the image will be blurred in the cross-scan  
15 direction. Also according to the ideal case, the angular scan speeds of the LOS due to the  
16 electronic scan and the optical rate and the scan rates must be also equal. If not, the pixels  
17 will be blurred in the along-scan direction. If both of these conditions are satisfied, then  
18 the outputs from the array will have the angular resolution as a single detector element, but  
19 an effective dwell time per pixel equal to N times the integration period of a single cycle,  
20 where N is the number of columns summed in the TDI (N=4 in the illustrated example).

21 Thus, in order to obtain the potential advantages of the TDI technique, the image  
22 rotation problem of the above-described prior art imaging system needs to be solved.

23 Another problem with the prior art has to do with how errors are inherently  
24 introduced across each scan line by variations in emissivity of the scan mirror. The GOES

1 imager measures its IR background by viewing deep space on one side of the Earth's  
2 surface at the end of each scan line. This background is subtracted from the raw scene  
3 measurements to determine the net radiance from the scene. Unfortunately, the emissivity  
4 of the GOES imager's scan mirror varies as a function of the angle of incidence and this  
5 angle of incidence varies by about  $9.5^\circ$  on each scan line. Because the background from  
6 the scan mirror is not constant during a scan line, the raw data exhibits east/west shading  
7 with errors of several degrees Kelvin in data from the Earth's surface. The error caused by  
8 this variation in emissivity over a single scan line is reduced, but not eliminated, by  
9 calibration. See M. Weinreb, et. al, "Operational calibration of Geostationary Operational  
10 Environmental Satellite-8 and -9 imagers and sounders", *Applied Optics*, vol. 36, no. 27,  
11 20 Sept. 1997, pp 6895-6904.

12 Referring to Fig. 5, the projection of a 2-D FPA into object space on alternate scan  
13 lines in a bi-directional scan is illustrated. The FPA is fabricated with a read-out column  
14 on each side of the active photo-detector array. Reversing the phasing of their charge-  
15 coupled transfer reverses the direction of motion of the charge packets in a row. These bi-  
16 directional TDI arrays permit bi-directional scanning, so the end of one scan line in an  
17 image can be in close angular proximity to the beginning of the next scan line. Between  
18 these two scan lines, both the opto-mechanical and electronic scan directions along the  
19 TDI axis must be reversed and the optical LOS must be offset in the cross-scan direction  
20 by an angle equal to (or slightly less than) the cross-scan angular subtense of the array.

21 There are other applications beside TDI that require that a sensor's opto-  
22 mechanical scan motion to lie along a fixed axis of a focal plane. For example, a  
23 multispectral scanner may have a focal plane that contains a series of linear FPA's,  
24 arranged as columns in the focal plane, each with a unique spectral filter. To achieve co-



1 registration among corresponding detector elements in several spectral channels, the same  
2 IFOV in the image must be scanned over the corresponding detector element in each FPA.

3 In prior scan mirrors on gimbal systems, including the GOES 8/9 imager, there is a  
4 variable relationship between the direction of scan produced by rotation about a single axis  
5 and the projection of the axes of the focal plane. In order to scan the image at a constant  
6 velocity (speed and direction) in the focal plane, it would be necessary to perform  
7 simultaneous, coordinated scanning at variable rates about the two gimbal axes. The outer  
8 gimbal must be capable of scanning over the full range of angles required to cover the  
9 FOR, continuously or in increments that are small in comparison to the IFOV of the  
10 sensor.

11 When the mirror is rotated about its inner axis, it changes the moment of inertia of  
12 the mass that is being rotated about the outer gimbal axis. Also, the outer gimbal axis  
13 cannot be a principal axis of the system for all mirror angles unless the mirror motion is  
14 compensated by a counter-rotating mass.

15 In most gimbal implementations, the outer gimbal must support the motor that  
16 drives the inner gimbal's rotation as well as an encoder or a resolver to make precision  
17 measurements of the inner gimbal's angle. Electric power and signals must be transferred  
18 cross the outer gimbal axis (e.g., by flexible cable or slip ring), increasing the rotational  
19 friction about the outer gimbal axis. Since highly accurate pointing is required, it is  
20 difficult to implement the necessary scan pattern by simultaneous rotations about both  
21 axes during the active portion of the scan.

22 Thus, what is needed is a scanning imager that avoids the problem of image  
23 rotation so that TDI may be effectively used. What is also needed is a scanning imager  
24 that has a reduced variation in emissivity of the scanning optics across each scan line.

1 What is also needed is a scanning mechanism that requires rotation about only one axis,  
2 preferably the inner axis, during the data-taking portion of the scan pattern.

### 3 SUMMARY OF THE INVENTION

4 The subject invention is a method and apparatus for scanning a two dimensional  
5 field of regard with a single plane mirror in the object space of a telescope, maintaining a  
6 fixed relationship between the rotational direction of scan and the projection of the  
7 telescope's focal plane. The two dimensional FOR is covered by a series of conical arcs,  
8 each arc being scanned by rotation at constant angular velocity about the inner axis of the  
9 two-axis system. This scanning system can accommodate applications such as TDI that  
10 require an opto-mechanical scan with a constant linear velocity (magnitude and direction)  
11 in the focal plane.

12 Shading of IR images is particularly hard to correct because the observed radiance  
13 has a component that varies with both the temperature of the scan mirror and its reflection  
14 angle. Calibration measurements, determined by viewing a source of known radiance, can  
15 mitigate this problem, but the time between calibration and scene measurements must be  
16 as short as possible or the calibration will be degraded by thermal drifts and 1/f noise in the  
17 detector and electronics.

18 In the subject invention, the reflection angle's dependence of the rotational angle  
19 reaches a minimum at the center of each arc in the FOR, and exhibits only a slow,  
20 quadratic increase towards each end of the arc as a the scan mirror is rotated about the  
21 inner gimbal's axis. The large variations in scan angle occur between arcs rather than  
22 within an arc. In the geosynchronous imager application, measurements of dark space,  
23 taken at the ends of the scan line, are used to subtract the instrument's background from  
24 the raw signals. The space measurements taken on each side of each arc may be used for

1 calibration of that arc, mitigating the shading problem. It is also an option to place one or  
2 more calibration sources, such as blackbodies with precision temperature-monitoring  
3 apparatus, in positions where they can be viewed between scan lines.

4 The only rotational motion during the active arc scanning is the rotation of the  
5 mirror at a constant angular velocity about its principal axis of minimum angular  
6 momentum. Because the outer axis of rotation remains fixed during the active portion of  
7 the scan, dynamic errors (jitter) in the outer gimbal's rotation and cross-coupling between  
8 the rotational motions about two axes are eliminated. Because the angular velocity is  
9 constant, torque disturbances due to angular acceleration about the inner axis are also  
10 eliminated during the active portion of the scan. All of these factors tend to reduce or  
11 eliminate errors in the pointing accuracy of the system.

12 ~~Between these arcs, the orientation of this inner axis is offset by rotation about the~~  
13 ~~outer gimbal axis. The outer gimbal needs only to be capable of holding a small number~~  
14 ~~of fixed positions, with the mechanical angle between positions no greater than one-half of~~  
15 ~~the optical width (cross-scan) of the scanned arcs. These factors tend to simplify the~~  
16 ~~apparatus required to measure and control the rotational angle of the outer axis. The~~  
17 ~~position of the outer gimbal must be known to the same level of accuracy as that of the~~  
18 ~~inner gimbal, however. The angular velocity of the mechanical scan is held constant~~  
19 ~~during any given arc. However is preferred to change the angular velocity from arc-to-arc,~~  
20 ~~in order to maintain the same optical scan rate for each arc.~~

21 It is an object of the present invention to provide a scanning imager that avoids the  
22 problem of image rotation so that the along-scan and cross-scan axes of the scan pattern  
23 correspond to fixed axes in the focal plane throughout the field of regard.

1 It is also an object of the present invention to provide a scanning imager that  
2 substantially eliminates image rotation so that TDI may be effectively used.

3 It is another object of the present invention to provide a scanning imager that has a  
4 reduced variation in emissivity of the scanning optics across each scan line.

5 It is yet another object of the present invention to provide a method for scanning an  
6 imager that compensates for the problem of image rotation.

7 It is still another object of the present invention to provide an imaging satellite that  
8 images a planet's surface while compensating for the problem of image rotation.

9 It is a further object of the present invention to provide a scanning imager that  
10 allows for effective co-registration among multiple detector arrays.

11 It is an additional object of the present invention to provide a scanning imager that  
12 permits multi-spectral imaging using multiple detector arrays.

13 It is another object of the present invention to provide a scanning imager that  
14 permits hyperspectral imaging.

15 It is also an object of the present invention to provide a scanning imager that  
16 allows the outer axis of a two-axis scanning gimbal to remain stationary during the data-  
17 taking portion of the scan pattern.

18 Some of the above objects are achieved by a method of scanning a field of view of  
19 an imager across a field of regard using a scan mirror mounted on a gimbal having an  
20 inner axis and an outer axis. The method includes sweeping the field of view across the  
21 field of regard in a selected direction by rotating the gimbal about the inner axis while  
22 maintaining the gimbal at a fixed angle with respect to the outer axis so as to. The method  
23 further includes progressing to a subsequent scan position by rotating the gimbal about the  
24 outer axis by a predetermined increment angle while maintaining the gimbal at a fixed

1 angle with respect the inner axis. Additionally, the method includes repeating the act of  
2 sweeping such that the selected direction is chosen alternately from a first direction and a  
3 second direction that is opposed to the first direction. The method further includes  
4 repeating the act of progressing prior to each repeated act of sweeping, wherein there is  
5 substantially no rotation, with respect to the instantaneous direction of scan, of an image  
6 formed on the imager.

7 Others of the above objects are achieved by an apparatus for scanning a two  
8 dimensional field of regard. The apparatus includes a telescope having a focal plane and a  
9 field of view, and one or more image sensors disposed at the focal plane. It also includes a  
10 single optically flat mirror disposed in the object space of the telescope, wherein the flat  
11 mirror scans the field of view across the field of regard while maintaining a fixed  
12 relationship between the rotational direction of scan and the projection of the telescope's  
13 focal plane.

14 Some of the above objects are obtained by an apparatus for imaging a two  
15 dimensional field of regard. The apparatus includes an imager having a field of view  
16 along a line of sight, the field of view being substantially smaller that the field of regard,  
17 as well as a scan mirror disposed so as to cast the line of sight onto the field of regard.  
18 The scan mirror causes the line of sight to be scanned across the field of regard in a  
19 conical arc.

## 20 BRIEF DESCRIPTION OF THE DRAWINGS

21 Additional objects and advantages of the present invention will be apparent in the  
22 following detailed description read in conjunction with the accompanying drawing figures.

23 Fig. 1 illustrates a schematic view of the geometric configuration of the prior art  
24 scanning optics for a GOES 8 & 9 imager.

1 Fig. 2 illustrates projection of flat panel arrays onto the surface of the Earth  
2 according to the prior art imager illustrated in Fig. 1.

3 Fig. 3 illustrates Time Delay and Integration, a prior art technique, showing the  
4 projection of a 4x4 flat panel array.

5 Fig. 4 illustrates Time Delay and Integration for a 4x4 array where the right-to-left  
6 opto-mechanical scan cancels the left-to-right electronic scan.

7 Fig. 5 illustrates the projection of a 2-D FPA into object space on alternate scan  
8 lines in a bi-directional scan.

9 Fig. 6 illustrates an opto-mechanical system according to an embodiment of the  
10 present invention for implementing a scan pattern across the full Earth disk from  
11 geostationary orbit.

12 Fig. 7 illustrates a bi-directional scan pattern for mapping the full Earth disk that is  
13 generated by the system of Fig. 6.

14 Fig. 8 illustrates a schematic view of the geometric configuration of scanning  
15 optics for the system of Fig. 6.

16 Fig. 9 illustrates projection of flat panel arrays onto the surface of the Earth  
17 according to the scanning optics illustrated in Fig. 8.

18 Fig. 10 illustrates pointing geometry for arc scanning according to the embodiment  
19 of Figs. 6-9.

20 Fig. 11 illustrates an opto-mechanical system according to a preferred embodiment  
21 of the present invention for implementing a scan pattern across the full Earth disk from  
22 geostationary orbit.

23 Fig. 12 illustrates a bi-directional scan pattern for mapping the full Earth disk that  
24 is generated by the system of Fig. 11.

1 Fig. 13 illustrates geometry to accommodate multiple linear FPA's covering  
2 several different spectral channels.

### 3 DETAILED DESCRIPTION OF THE INVENTION

4 The scanning technique proposed here mitigates the three problems discussed  
5 above by scanning a series of arcs across the FOR so that the opto-mechanical scan  
6 velocity (magnitude and direction) as projected into the focal plane, remains constant  
7 during the active scan. Each arc is scanned by rotation at constant angular velocity about  
8 the inner axis of a two-axis gimbal system while the outer gimbal axis remains fixed.  
9 Between arcs, the direction of rotation about the inner axis is reversed and the angle of the  
10 outer axis is offset. Reference measurements of dark space may also be taken during this  
11 turn-around interval, at reflection angles that are very close to the reflection angles of the  
12 preceding and following arcs. Alternatively, one or more calibration sources, such as  
13 infrared blackbodies with precision temperature measurement, may also be placed in  
14 positions where they do not vignette the FOR, but where they can be viewed during the  
15 interval between consecutive arcs.

16 Referring to Fig. 6, an opto-mechanical system according to an embodiment of the  
17 present invention for implementing a scan pattern across the full Earth disk from  
18 geostationary orbit is illustrated. The primary mirror **612** and secondary mirror **614** form a  
19 telescope **610** with its line-of-sight (LOS) pointing south, in the +y direction,  $\mathbf{I} = [0, +1,$   
20  $0]$ . The along-scan axis of the FPA **620** is parallel to the E/W axis of the spacecraft (x-  
21 axis) and the cross-scan axis is parallel to the zenith/nadir axis (z-axis).

22 The scan mirror **630** is an optical flat with an elliptical cross-section that is  
23 mounted on a gimbal system **640** with two perpendicular axles **642**, **644**. The outer axle  
24 **642** is attached to the body **602** of the instrument parallel to its x-axis and perpendicular to

1 the LOS of the telescope. The projection of the cross-scan axis of the FPA **620** onto the  
2 scan mirror **630** is always perpendicular to the outer axle **642** of the gimbal. The outer  
3 gimbal axle **642** is parallel to the TDI axis of the FPA **620**. The inner axle **644** of the  
4 gimbal **640** rotates about the x-axis of the spacecraft due to rotation about the outer axle  
5 **642** and assumes different orientations in the y/z plane. The scan mirror **630** is attached to  
6 the inner axle **644** with the normal vector of its mirrored surface perpendicular to the inner  
7 axle **644**. The major axis of its elliptical cross-section is parallel to the inner axle **644**.

8 In its home position, the scan mirror's normal vector points 45 degrees north and  
9 45 degrees towards nadir,  $\mathbf{N} = [0, -0.707, +0.707]$ , and the telescope's line of sight,  $\mathbf{I} = [0,$   
10  $+1, 0]$ , is reflected toward the center of the Earth's disk, at nadir,  $\mathbf{R} = [0, 0, +1]$ .

11 An arc is scanned across the Earth's disk by holding the outer axle **642** of the  
12 gimbal **640** in a fixed position and rotating the scan mirror **630** about the inner axle **644** at  
13 a constant angular velocity. At the end of a scan line, the scan mirror/inner axle assembly  
14 is rotated about the outer axle **642** while the direction of rotation about the inner axis is  
15 reversed. The direction of TDI motion on the FPA **620** is also reversed.

16 Referring to Fig. 7, a bi-directional scan pattern **710** for mapping the full Earth  
17 disk **720** that is generated by the system of Fig. 6 is illustrated. The scan pattern **710**  
18 consists of alternating west-to-east and east-to-west arcs, proceeding from north-to-south,  
19 with an angle of 0.75 degrees between consecutive arcs, on the central meridian **722**,  
20 between the North Pole **724** and South Pole **726**. Projections **716** of FPA pixels at various  
21 scan positions are shown. The nominal cross-scan dimension of the instrument's FOV is  
22 0.8 degrees in this example. Thus, this scan pattern **710** produces some overlap between  
23 adjacent arcs.



1 To scan the first arc **712** in this pattern, the scan mirror/inner axle assembly is  
2 slewed away from its home position by rotating 4.5 degrees counterclockwise (refer to  
3 Fig. 6) about the outer axle **642**, so that the scan mirror's normal vector points 49.5  
4 degrees northward and the angle of reflection of the telescope's LOS from the mirror is  
5 40.5 degrees. The mirror **630** is then rotated approximately 6 degrees about the inner axle  
6 **644** so that it points 9 degrees west of the central meridian. The outer axle **642** is held  
7 stationary and the mirror **630** is rotated about the inner axle **644** of the gimbal **640** at a  
8 constant angular velocity, causing the reflected LOS to scan from west-to-east. The image  
9 of the FPA rotates so that the projection of its TDI axis remains along the direction of  
10 motion.

11 At the end of the first arc **712**, scan mirror/inner axle assembly is rotated about the  
12 outer axle **642** by 0.375 degrees clockwise about the x-axis, causing the reflected LOS of  
13 the second arc **714** to be 0.75 degrees further south than the first arc **712** on the central  
14 meridian **722**. The direction of rotation about the inner axle **644** is reversed, causing the  
15 second arc **714** to scan from east-to-west. At the end of each scan line, the angle of outer  
16 axle **642** is incremented by 0.375 degrees clockwise and the direction of rotation about the  
17 inner axle **644** is reversed. Fig. 7 shows a sequence of 25 arcs covering the full-Earth disk  
18 **720** with a series of "frowns." The angular velocity of the LOS is predominately eastward  
19 or westward, but it initially has a small northward component that vanishes when the arc  
20 crosses the central meridian and then turns southward in the second half of the arc.

21 Reference measurements are made at the beginning and the end of each arc, while  
22 the LOS is pointing to space. The reflection angle from the scan mirror reaches a  
23 minimum at the center of each arc in the FOR. It exhibits only a slow, quadratic increase  
24 towards each end of the arc (the eastern and western sides of the Earth's disk) as the scan

1 mirror is rotated about the inner gimbal's axis. The large variations in scan angle occur  
2 between arcs rather than within an arc. In this geosynchronous imager embodiment,  
3 measurements of dark space are used as references to subtract the instrument's background  
4 from the raw signals. The time interval between reference and scene measurements is  
5 short, minimizing degradations caused by thermal drifts and 1/f noise.

6 The only rotational motion during the active arc scanning is the rotation of the  
7 mirror at a constant angular velocity. Because the outer axis of rotation remains fixed  
8 during the active portion of the scan, dynamic errors (jitter) in the outer gimbal's rotation  
9 and cross-coupling between the rotational motions about two axes are eliminated.  
10 Because the angular velocity is constant, torque disturbances due to angular acceleration  
11 about the inner axis are also eliminated during the active portion of the scan. All of these  
12 factors tend to reduce or eliminate errors in the pointing accuracy of the system.

13 The outer axis of the gimbal 642 needs only to be capable of holding a small  
14 number of fixed positions, with the mechanical angle between positions no greater than  
15 one-half of the cross-scan angle of the FOV of the optical system. These factors tend to  
16 simplify the apparatus required to measure and control the rotational angle of the outer  
17 axis. It is preferable that the position of the outer axis of the gimbal be known to the same  
18 level of accuracy as that of the inner axis of the gimbal.

### 19 Optical Geometry

20 Referring to Fig. 8, a schematic view of the optical geometry for the embodiment  
21 of Fig. 6 is illustrated. The coordinate axes are: +x = east, +y = south, +z = nadir. The  
22 telescope's optical axis 810 points south, in the +y direction. Its focal plane 812 coincides  
23 with the x-z plane of the spacecraft. The TDI axis 814 of the FPA, indicated by the arrow  
24 in the focal plane, is aligned to the x-axis. The cross-scan axis 816 of the FPA, indicated

1 by the "lollypop," is aligned to the z-axis of the focal plane. Optionally, the satellite  
2 performs a yaw flip at each equinox, so the telescope points north, the +x-axis points west  
3 and the +y north for six months of the year. At nadir, the cross-scan axis 816 of the FPA  
4 is projected in the north/south direction and the TDI axis 814 is projected in the east/west  
5 direction.

6 The optical geometry of Fig. 8 may be compared with the prior art illustrated in  
7 Fig. 1. Both figures use the same coordinate convention, +x = east, +y = south, +z =  
8 nadir. In both figures, the outer axis of the gimbal is parallel to the x-axis of the inner axis  
9 of the gimbal has a variable orientation perpendicular to the x-axis. Both systems scan a  
10 FOR that includes the Earth's surface as viewed from geostationary orbit by scanning  
11 alternate lines from east-to-west and west-to-east and by stepping from north-to-south  
12 between scan lines. In both cases, the scan lines are implemented by rotation of the scan  
13 mirror about its inner axis and the north to south steps are implemented by rotation of the  
14 yoke about the outer axis of the gimbal system during the time interval between scan lines.

15 The orientations of the telescopes differ between the two figures. In Fig. 1, the  
16 optical axis of the telescope points east, parallel to the +x axis, but in Fig. 8, the optical  
17 axis of the telescope points south, parallel to the +y axis. In both cases, the angle of  
18 reflection of the telescope's axis from the scan mirror is 45 degrees when the FPA is  
19 projected toward nadir. In both cases, the scan mirror is an optical plane mirror with an  
20 elliptical cross-section. In Fig. 1, the inner axis of rotation coincides with the minor axis  
21 of the ellipse. In Fig. 8, the inner axis of rotation coincides with the major axis of the  
22 ellipse.

23 Fig. 8 may also be compared with Fig. 6. These two figures show different  
24 embodiments of the scan mirror's gimbal system. Whereas the gimbal system of Fig. 6 is

fixed to the backside of the scan mirror, the gimbal system of Fig. 8 uses a yoke around the scan mirror. The embodiment of Fig. 8 allows the center of rotation (the intersection of the two gimbal axes) to lie on or close to the surface of the mirror. This condition is desirable to minimize the size of the scan mirror and to use the same area on the scan mirror for all measurements. The embodiment of Fig. 6 eliminates the bulky yoke and its accompanying inertia. The orientation of the axes of the gimbal with respect to the x-y-z axis system of the spacecraft, and with respect to the telescope, is the same in both embodiments.

Referring to Fig. 9, projections **902, 904, 906, 912, 914, 916** of the crossed axes in the focal plane onto the Earth's surface **30** are illustrated. The image rotation as is illustrated by the TDI axes of each projection is compensated by the geometry of the arcs so that the along scan axis of the focal plane remains aligned with the direction of the scan over the entire FOR. This is in stark contrast to Fig. 2 which illustrates the mis-alignment between the scan direction and the focal plane array due to image rotation in the scan line across the North Pole according to the prior art. Figs. 2 and 9 both illustrate FPA projections of two scan lines: one at the Equator, and another near the North Pole.

The curvature of the scanning arcs has been exaggerated in Fig. 9. The actual scan pattern is illustrated in Fig. 7.

### **Analysis of Arc Scanning Motion**

Referring to Fig. 10, vectors **I**, **N**, and **R** relevant to the pointing geometry for arc scanning according to the embodiment of Figs. 6-9 are illustrated. The following analysis explains how the scan is coordinated with TDI in the focal plane.

The entire focal plane, with several FPA's that scan in sequence along an arc, may also subtend a small angle in each direction. In this analysis, the field-of-view (FOV) of

the focal plane is assumed to be square and to extend  $\pm 0.4^\circ$  ( $\pm 0.007$  radians) in each direction from its center to its corners. The telescope projects the center of the focal plane, ( $x=0, z=0$ ) along the y-axis. If the EFL of the telescope is  $f$ , then the telescope projects the line of sight (LOS) from each point ( $x,z$ ) in the focal plane in the direction defined by the angles:

$$A = \arcsin(z/f), B = \arcsin(x/f) \quad [1]$$

The telescope collimates the line of sight (LOS) from each point in the focal plane into a direction  $\mathbf{I(A,B)}$ . In this analysis, vector quantities are typeset as boldfaced; their components are enclosed in bold brackets. The center of the focal plane, for which  $A=0, B=0$ , is projected onto the scan mirror along the direction defined by:

$$\mathbf{I(A=0,B=0)} = [0,1,0] \quad [2]$$

Third and higher order terms in  $A$  and  $B$  are no greater than  $(0.007 \text{ radian})^3 = 3.4 \times 10^{-7}$  radian and are neglected in this analysis. In this approximation, the telescope projects an arbitrary point ( $x,z$ ) in the focal plane into a LOS,  $\mathbf{I(A,B)}$  given by:

$$\mathbf{I(A,B)} = [B, 1 - \frac{1}{2}(A^2+B^2), A] \quad [3]$$

The scanning system has a plane scan mirror **818** with an elliptical cross-section mounted on a two-axis gimbal system. The inner gimbal axis **820** coincides with the long axis of the ellipse and is perpendicular to the mirror's normal vector  $\mathbf{N}$ . The outer gimbal axis **824** is perpendicular to the inner axis **820** and parallel to the east/west axis ( $x$ -axis) of the spacecraft. Rotation about the outer gimbal axis **824** causes the orientation of the inner gimbal axis **820** to rotate in the  $y/z$  plane. The outer gimbal angle ( $\text{oga}$ ) is defined to be the angle between the  $y$ -axis and inner gimbal axis **820**. Rotation about the inner gimbal axis **820** causes the orientation of the scan mirror's normal vector  $\mathbf{N}$ , to rotate in the plane perpendicular to the outer gimbal axis **824**. The inner gimbal angle ( $\text{iga}$ ) is defined to be

the angle between the scan mirror's normal vector  $\mathbf{N}$  and the y/z plane. The normal vector is, therefore, a function of the two gimbal angles,  $\mathbf{N}(\text{iga}, \text{oga})$ . The rays from an arbitrary point in the focal plane are collimated into a LOS,  $\mathbf{I}(\mathbf{A}, \mathbf{B})$ , that is a function of the two field angles, A and B. The reflected LOS,  $\mathbf{R}(\mathbf{A}, \mathbf{B}, \text{iga}, \text{oga})$ , is a function of the two field angles, A and B, and the two gimbal angle, iga and oga. The incident vector  $\mathbf{I}$ , the reflected vector  $\mathbf{R}$ , and the mirror's normal vector  $\mathbf{N}$ , are related by the standard equation:

$$\mathbf{R}(\mathbf{A}, \mathbf{B}, \text{iga}, \text{oga}) = \mathbf{I}(\mathbf{A}, \mathbf{B}) - 2\mathbf{N}(\text{iga}, \text{oga})\{\mathbf{N}(\text{iga}, \text{oga}) \cdot \mathbf{I}(\mathbf{A}, \mathbf{B})\} \quad [4]$$

where  $\mathbf{N}(\text{iga}, \text{oga}) \cdot \mathbf{I}(\mathbf{A}, \mathbf{B})$  is the scalar product of the two vectors.

To find the general equation for  $\mathbf{N}(\text{iga}, \text{oga})$ , consider the state in which  $\text{oga} = 0$ , where the outer gimbal axis coincides with the y-axis (the optical axis of the telescope) and  $\text{iga} = 0$ , so the normal vector lies in the y/z plane. In this position, the normal vector  $\mathbf{N}$  points toward nadir:

$$\mathbf{N}(0,0) = [0,0,1] \quad [5]$$

The scan mirror **818** may be rotated about the inner axis so that it assumes an arbitrary orientation in the x/z plane. The normal vector  $\mathbf{N}$  as a function of iga is now:

$$\mathbf{N}(\text{iga}, 0) = [\sin(\text{iga}), 0, \cos(\text{iga})] = [S_i, 0, C_i] \quad [6]$$

where  $S_i = \sin(\text{iga})$  and  $C_i = \cos(\text{iga})$ .

The home position of the scan mirror **818** may be defined as that position that reflects the line of sight (LOS) from the center of the FPA,  $\mathbf{I}(\mathbf{A}=0, \mathbf{B}=0) = [0,1,0]$  to nadir, at  $\mathbf{R}(\mathbf{A}=0, \mathbf{B}=0, \text{iga}, \text{oga}) = [0,0,1]$ , in the center of the field of regard (FOR) that the system is required to scan. From Eq. [4], it is straightforward to show that the home position of the gimbal system is reached when  $\text{iga} = 0^\circ$  and  $\text{oga} = +45^\circ$ , so that:

$$\mathbf{N}(\text{iga}=0, \text{oga}=45) = [0, -0.707, +0.707] \quad [7]$$

The general equation for  $N(\text{iga}, \text{oga})$ , is determined by rotation of  $N(\text{iga}, 0)$  as defined by Eq [6] about the oga, which coincides with the x-axis of the spacecraft:

$$N(\text{iga}, \text{oga}) = [S_i, B \sin(\text{oga})C_i, \cos(\text{oga})C_i] \\ = [S_i, B S_o C_i, C_o C_i] \quad [8]$$

where  $S_o = \sin(\text{oga})$  and  $C_o = \cos(\text{oga})$ .

The general case for a reflected angle,  $R(A, B, \text{iga}, \text{oga})$ , from any position (A,B) in the FPA with any orientation of  $N(\text{iga}, \text{oga})$ , is determined by combining Eqs [3], [4], and [8]:

$$R(A, B, \text{iga}, \text{oga}) = [B, 1 - \frac{1}{2}(A^2 + B^2), A] \\ - 2[S_i, -S_o C_i, C_o C_i] \{ [S_i, -S_o C_i, C_o C_i] \cdot [B, 1 - \frac{1}{2}(A^2 + B^2), A] \} \quad [9]$$

When third and higher order terms in A and B are neglected, Eq. [9] reduces to:

$$R_x(A, B, \text{iga}, \text{oga}) = 2S_o S_i C_i - 2A C_o S_i C_i + B(1 - 2S_i^2) - (A^2 + B^2)S_o S_i C_i \\ R_y(A, B, \text{iga}, \text{oga}) = (1 - 2S_o^2 C_i^2) + 2A S_o C_o C_i^2 \\ + 2B S_o S_i C_i - \frac{1}{2}(A^2 + B^2)(1 - 2S_o^2 C_i^2) \\ R_z(A, B, \text{iga}, \text{oga}) = 2S_o C_o C_i^2 + A(1 - 2C_o^2 C_i^2) \\ - 2B C_o S_i C_i - (A^2 + B^2)S_o C_o C_i^2 \quad [10]$$

### Coordination of Electronic and Opto-Mechanical Scanning

The projection of the TDI axis of the FPA onto the Earth at any gimbal angle may be found by differentiating Eq.[10] with respect to B. To simplify, apply double-angle trigonometric identities to iga:

$$dR(A, B, \text{iga}, \text{oga})/dB = [1 - 2S_i^2 - 2B S_o S_i C_i, 2S_o S_i C_i - B(1 - 2S_o^2 C_i^2), \\ - 2C_o S_i C_i - 2B S_o C_o C_i^2] \\ = [\cos(2\text{iga}) - B S_o \sin(2\text{iga}), S_o \sin(2\text{iga}) \\ - B(1 - 2S_o^2 C_i^2), - C_o \sin(2\text{iga}) - 2B S_o C_o C_i^2] \quad [11]$$

Consider an array that is centered on the optical axis in the along-track direction (x-axis) of the focal plane, for which  $B=0$ . TDI transfers the charge packet in each row detectors in the FPA at an effective angular rate,  $dB/dt$ , equal to the linear rate of charge transfer,  $dx/dt$ , divided by the effective focal length of the telescope,  $f$ . If the gimbal axes

are held stationary, then the projected image of a charge packet has the angular velocity,

$V_{\text{charge}}$ :

$$V_{\text{charge}} = (dB/dt)dR(A,0,iga,oga)/dB \\ = (dB/dt)[\cos(2iga), S_o \sin(2iga), -C_o \sin(2iga)] \quad [12]$$

The projected line of sight from an array that is centered in the cross-scan direction

(B=0) is:

$$R_x(A,0,iga,oga) = 2S_o S_i C_i - 2A C_o S_i C_i - A^2 S_o S_i C_i \\ R_y(A,0,iga,oga) = (1 - 2S_o^2 C_i^2) + 2A S_o C_o C_i^2 - \frac{1}{2}A^2(1 - 2S_o^2 C_i^2) \\ R_z(A,0,iga,oga) = 2S_o C_o C_i^2 + A(1 - 2C_o^2 C_i^2) - A^2 S_o C_o C_i^2 \quad [13]$$

If the inner gimbal axis is rotated at an angular velocity of  $w$ , then:

$$dS_i/dt = w C_i \text{ and } dC_i/dt = -w S_i \quad [14]$$

Therefore, if the outer gimbal angle,  $oga$ , remains constant while the inner gimbal angle,  $iga$ , is rotated at a constant angular velocity,  $w$ , the line of sight of the optical axis changes at the rate:

$$dR(A,0,iga,oga)/dt = \{dR(A,0,iga,oga)/d(iga)\} \{d(iga)/dt\} \\ = w \{dR(A,0,iga,oga)/d(iga)\} \quad [15]$$

Combining Eqs. [13], [14], and [15] and applying the double-angle identities yields

the result:

$$dR_x(A,0,iga,oga)/dt = 2w \cos(2iga)[S_o - A C_o - \frac{1}{2}A^2 S_o] \\ dR_y(A,0,iga,oga)/dt = 2w \sin(2iga)[S_o^2 - A S_o C_o - \frac{1}{2}A^2 S_o^2] \\ dR_z(A,0,iga,oga)/dt = -2w \sin(2iga)[S_o C_o - A C_o^2 - \frac{1}{2}A^2 S_o C_o] \quad [16]$$

The angular velocity of the projection of a moving charge packet in the center of the focal plane, where  $A=0$  and  $B=0$  can be canceled by the opto-mechanical scan velocity when:

$$V_{\text{charge}} + dR(0,0,iga,oga)/dt = 0 \quad [17]$$

This is the condition where the electronic and opto-mechanical scan rates of the charge packet cancel in object space. It is possible to achieve this cancellation throughout



the two-dimensional field of regard because  $V_{\text{charge}}$  and  $d\mathbf{R}(0,0,\text{iga},\text{oga})/dt$  remain parallel for all values of iga and oga. Combining Eqs. [12], [16], and [17] yields the result:

$$\{dB/dt + 2wS_o\}[\cos(2\text{iga}), S_o\sin(2\text{iga}), -C_o\sin(2\text{iga})] = 0 \quad [18]$$

Therefore, for each scanning arc, the angular velocity,  $w$ , of the inner axis should be set equal to the following value, a function of the outer gimbal angle, oga:

$$w = -(dB/dt)/(2S_o) = -(dx/dt)/(2fS_o) \quad [19]$$

Note that the value of  $w$  remains constant for each arc but varies from arc to arc as a function of the outer gimbal angle, oga.

### Errors Due To Velocity Mis-Match in Cross-Scan Direction

When a linear FPA is centered in the along-scan direction,  $B=0$ , but  $A$  assumes values that span the field of view (FOV). According to this embodiment,  $A$  ranges between  $\pm 0.007$  radians, corresponding to  $\pm 0.4^\circ$  in object space and  $\pm 250$  km on the Earth's surface. When  $A$  is non-zero, there is a slight mis-match between the electronic and opto-mechanical scan rates. The opto-mechanical rate is too fast at the outer radius of the scanning arc (on the northern side, as shown in Fig. 7) and too slow at the inner radius. The velocity of the LOS vector, as a function of  $A$ ,  $d[\mathbf{R}(A,0,\text{iga},\text{oga})]/dt$ , may be compared to its value at the center of the array by using Eq. [16].

$$d\mathbf{R}(A,0,\text{iga},\text{oga})/dt = [1 - A\cot(\text{oga}) - 2A^2] d\mathbf{R}(0,0,\text{iga},\text{oga})/dt \quad [20]$$

The worst-case occurs for detectors elements at the end of the FPA, where  $A = +0.0070$ , when the AGSI is scanning the South Pole, for which  $\text{oga} = 45^\circ - \frac{1}{2}(8.7^\circ) = 40.65^\circ$ . In this worst case, the electronic and opto-mechanical scan rates are mis-matched by a scan rate error of less than 1%:

$$\begin{aligned} \text{Worst-case scan rate error} &= Aw\cot(40.65^\circ) + \frac{1}{2}A^2 \\ &= [(0.0070)(1.165) + \frac{1}{2}(0.0070)^2]w = +0.0082w \end{aligned} \quad [21]$$

1 The number of contiguous pixels integrated in the TDI process varies from channel  
2 to channel. For this embodiment of the present invention, the largest value of N, the  
3 number of pixels to be integrated by TDI, is N=20. The worst-case value of RMS blur due  
4 to this velocity error, expressed as a fraction of a pixel's dimension, is:

5 
$$\text{Worst-case RMS along-scan blur} = (0.0082)(20)/(12^{1/2})$$
  
6 
$$= 0.047 \text{ pixel} \quad [22]$$

7 Since the RMS blur induced by this mis-match is less than 5% of one pixel, it is a  
8 small source of image quality degradation.

### 9 Variation of Reflection Angle in a Single Scan Line

10 In the thermal IR channels, each mirror in the optical train emits background  
11 radiation that contributes to all signals measured by the detectors. This background from  
12 each optical element is proportional to its emissivity and increases with temperature  
13 according to Planck's law. Space looks, taken at the ends of each scan line, are used to  
14 measure the dark current due to instrumental background. This background is then  
15 subtracted from the raw scene data to determine the net radiance from the Earth. All  
16 components of the background due to thermal radiation from within the instrument are  
17 properly subtracted from the signal whenever they remain constant during a single scan  
18 line. All optical elements have temperatures that remain virtually constant during a single  
19 scan line, and all except the scan mirror 818 reflect the radiation from the scene in a fixed  
20 orientation, so their emissivities also remain constant. The scan mirror 818, however,  
21 reflects radiation at a variable angle over the course of a scan line. When the scan mirror's  
22 emissivity varies as a function of the reflection angle, it creates a bias in this background  
23 subtraction process that is proportional to the difference in emissivity between the  
24 reflection angle of the scene measurement and that of the background measurement.

5/15/12  
AS  
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Therefore, variation in the reflection angle within a scan line is more troublesome than variation in the reflection angle from line to line.

The reflection angle of the optical axis of the imager according to this embodiment, for which  $A=0$  and  $B=0$ , can be determined from:

$$\text{Reflection angle} = \arccos\{N(\text{iga}, \text{oga}) \cdot I(0,0)\} = \arccos(S_o C_i) \quad [23]$$

The arc that passes through the Equator is defined by  $\text{oga}=45^\circ$ . Along this arc,  $\text{iga}$  varies from  $\text{iga}=0^\circ$  at the center of this arc to  $\text{iga} = \pm 8.7^\circ/2^{1/2} = \pm 6.2^\circ$  at the Earth's limb.

The variation of the reflection angle between the center of this single arc through the Equator on the Earth's surface is:

$$\begin{aligned} \text{Reflection angle variation} &= \arccos\{\sin(45^\circ)\cos(0^\circ)\} \\ &\quad - \arccos\{\sin(45^\circ)\cos(6.2^\circ)\} \\ &= 45^\circ - 45.3^\circ = -0.3^\circ \end{aligned} \quad [24]$$

In contrast, the reflection angle of the prior art GOES Imager's scan mirror varies by  $8.7^\circ$ , from  $40.65^\circ$  at the western limb to  $49.35^\circ$  at the eastern limb, in a single west-to-east scan line across the Equator. The present invention's reflection angle increases by  $8.7^\circ$  from North Pole to South Pole, but each successive arc across the Earth's surface is bracketed by space-viewing measurements that measure the dark current at nearly the same scan mirror reflection angle. Therefore, the shading that is present in the uncalibrated prior art GOES data is reduced by at least an order of magnitude by the alternative scan technique according to the present invention.

### Optimal Alignment of Multiple FPA's

Preferably, an imager according to the present invention accommodates multiple linear FPA's covering several different spectral channels in each of its four focal planes. The focal planes are extended in the along-scan direction, so that FPA's can have values of

1 B ranging from -0.0070 to +0.0070. Each FPA is properly oriented with its TDI axis  
2 along the direction of the opto-mechanical scan.

3 Rotation about the inner gimbal axis projects a fixed line of sight from the  
4 telescope into a conical arc across the Earth's surface. Due to the reversibility of optical  
5 systems, this rotation also projects the line of sight from a fixed point on the Earth's  
6 surface into a conical arc in the focal plane of the telescope. Therefore, the best co-  
7 registration among the linear FPA's is achieved when they are aligned with their cross-  
8 scan axes pointing toward the axis of the cone, like the spokes of a wheel. This geometry  
9 is illustrated in Fig. 13.

10 The projections of all of the FPA's, A through E, are aligned with their cross-scan  
11 axes converging on the extension of the inner gimbal axis. When the  $\text{iga} = 0^\circ$ , the LOS  
12 from the each array is reflected into the corresponding FOV, A' through E'. When the  $\text{iga}$   
13 is rotated in the positive direction, the central array, C, is scanned from C' to B' to A',  
14 along an arc as previously described. In this radial focal plane alignment, all of the other  
15 FPA's will also scan in sequence through the same arc with their TDI axes along the  
16 direction of scan.

17 The optimal FPA alignment varies as a function of the  $\text{oga}$ . The following analysis  
18 assumes that this alignment is optimized at  $\text{oga} = 45^\circ$ , the arc that scans the Equator. In  
19 this configuration, all of the FPA's are properly aligned when scanning through the  
20 Equator, but they become progressively more misaligned at higher latitudes, reaching their  
21 worst case mis-alignment at the poles, where  $\text{oga} = 45^\circ \pm 4.35^\circ$ .

22 Consider a FPA that is centered in the along-scan direction, so that  $B=0$ . The  
23 telescope projects this FPA into a LOS,  $I(A,0)$ , where A ranges from -0.0070 to +0.0070  
24 to cover the cross-scan width of the array.  $I(A,0)$  is determined by setting  $B=0$  in Eq. [3]:

$$\mathbf{I}(\mathbf{A}, \mathbf{0}) = [0, 1 - \frac{1}{2}A^2, A] \quad [25]$$

When  $\text{iga} = 0$ , the scan mirror's normal vector, determined from Eq. [6], is:

$$\mathbf{N}(\mathbf{0}, \text{oga}) = [0, BS_o, C_o] \quad [26]$$

The reflection of  $\mathbf{I}(\mathbf{A}, \mathbf{0})$  from the scan mirror into object space is:

$$\mathbf{R}(\mathbf{A}, \mathbf{0}, \mathbf{0}, \text{oga}) = [0, 1 - \frac{1}{2}A^2, A] - 2[0, -S_o, C_o]\{[0, -S_o, C_o] * [0, 1 - \frac{1}{2}A^2, A]\} \quad [27]$$

The components of Eq. [27] are:

$$\begin{aligned} R_x(\mathbf{A}, \mathbf{0}, \mathbf{0}, \text{oga}) &= 0 \\ R_y(\mathbf{A}, \mathbf{0}, \mathbf{0}, \text{oga}) &= (1 - 2S_o^2) + 2AS_oC_o - \frac{1}{2}A^2(1 - 2S_o^2) \\ R_z(\mathbf{A}, \mathbf{0}, \mathbf{0}, \text{oga}) &= 2S_oC_o + A(1 - 2C_o^2) - A^2S_oC_o \end{aligned} \quad [28]$$

We now wish to determine the proper alignment for a FPA displaced by an angle B in the along scan direction. An angular displacement of B corresponds to a rotation about the inner gimbal axis from  $\text{iga} = 0$  to:

$$\text{iga} = -B/(2S_o) \quad [29]$$

Combining Eqs. [8] and [29] and using the second-order approximations for  $S_i$  and  $C_i$ , the scan mirror's normal vector becomes:

$$\mathbf{N}(-B/(2S_o), \text{oga}) = [-B/(2S_o), -S_o + B^2/(8S_o), C_o - B^2C_o/(8S_o^2)] \quad [30]$$

The incident LOS that is projected onto the same footprint on the Earth's surface may be found by combining Eqs. [28] and [30] with the inverse of Eq. [4]:

$$\mathbf{I}(\mathbf{A}, \mathbf{B}, -B/(2S_o), \text{oga}) = \mathbf{R}(\mathbf{A}, \mathbf{0}, \mathbf{0}, \text{oga}) - 2\mathbf{N}(-B/(2S_o), \text{oga})\{\mathbf{N}(-B/(2S_o), \text{oga}) * \mathbf{R}(\mathbf{A}, \mathbf{0}, \mathbf{0}, \text{oga})\} \quad [31]$$

Neglecting third and higher order terms in A and B, this equation becomes:

$$\begin{aligned} I_x(\mathbf{A}, \mathbf{B}, -B/(2S_o), \text{oga}) &= B - ABC_o/S_o = B - AB\cot(\text{oga}) \\ I_y(\mathbf{A}, \mathbf{B}, -B/(2S_o), \text{oga}) &= 1 - \frac{1}{2}(A^2 + B^2) \\ I_z(\mathbf{A}, \mathbf{B}, -B/(2S_o), \text{oga}) &= A + 2B^2C_o/S_o = A + \frac{1}{2}B^2\cot(\text{oga}) \end{aligned} \quad [32]$$

The angular alignment for the cross-scan axis of the array in the focal plane (the x-z plane) can be found by differentiating  $I_x$  and  $I_z$  with respect to the cross-scan variable, A.

$$1 \quad \text{Optimal tilt angle} = [d I_x/dA]/[d I_z/dA] = -B \cot(\text{oga}) \quad [33]$$

2 Therefore, optimal tilt angle between each FPA and the z-axis, in radians, equals  
 3  $B \cot(\text{oga})$ . To scan each FPA over the same area on the Earth's surface, the center of each  
 4 array ( $A=0$ ) should be located at:

$$5 \quad \text{Optimum array center} = [I_x = B, I_z = \frac{1}{2}B^2 \cot(\text{oga})] \quad [34]$$

6 Each FPA should be tilted to satisfy Eq [33] and offset to satisfy Eq [34] on the  
 7 central arc, that passes through nadir and scans the Equator, for which  $\text{oga} = 45^\circ$ . If the  
 8 FPA's are all parallel to one another (and to the z-axis), then the maximum tilt error occurs  
 9 for FPA's at the edges of the focal plane, where  $B = \pm 0.007$  radian, and for  $\text{oga} = 40.65^\circ$ ,  
 10 the arc that scans through the South Pole. This worst-case tilt angle error, in radians, is:

$$11 \quad \text{Worst-case tilt error (parallel FPA's)} = \\ 12 \quad \pm B \cot(40.65) = \pm 0.0070(1.165) = \pm 0.0082 \quad [35]$$

13 When TDI is implemented over 20 consecutive pixels, the worst-case RMS cross-  
 14 scan blur, expressed as a fraction of a pixel, is:

$$15 \quad \text{Worst-case cross-scan blur (parallel FPA's)} = \\ 16 \quad (0.0082)(20)/(12^{1/2}) = 0.047 \text{ pixel} \quad [36]$$

17 This cross scan blur associated with parallel FPA's is probably tolerable, but it can  
 18 be reduced by tilting the arrays to satisfy Eq [33]. With this radial focal plane alignment,  
 19 the maximum tilt error also occurs at the South Pole, but it is reduced to

$$20 \quad \text{Worst-case tilt error (radial FPA's)} = \pm B[\cot(40.65)-1] \\ 21 \quad = \pm 0.0070(1.165-1) \\ 22 \quad = \pm 0.00116 \quad [37]$$

23 The RMS cross-scan blur, expressed as a fraction of a pixel, becomes:

$$24 \quad \text{RMS cross-scan blur (radial FPA's)} = (0.00116)(20)/(12^{1/2}) \\ 25 \quad = 0.0067 \text{ pixel} \quad [38]$$

1 This cross scan blur, less than 1% of the pixel's dimension, is negligible when the FPA's  
 2 are aligned radially. Comparison of Eq [36] against Eq [38] shows that the worst-case  
 3 blur circle due to the off-axis FPA's tilt error is about seven times smaller if they are  
 4 aligned radially rather than in parallel.

5 A preferred feature of an imager according to the present invention is channel-to-  
 6 channel co-registration i.e., to have corresponding pixels in its different spectral bands of  
 7 its processed imagery measure the radiance from same area on the Earth's surface. The  
 8 present invention imager achieves co-registration in its processed imagery by scanning  
 9 with detector elements that subtend smaller instantaneous fields of view (IFOV's) that  
 10 those of the pixels in the image, and then re-sampling to the image grid registered to the  
 11 Earth's surface. Although this approach allows the hardware co-registration requirements  
 12 to be relaxed, it is still desirable to have corresponding pixels in the different FPA's scan  
 13 the IFOV to the greatest extent possible.

14 One aspect of hardware co-registration is that FPA's are aligned in the cross-scan  
 15 direction so that the centers of corresponding detector elements in each array are scanned  
 16 over the same IFOV in object space. If all of the FPA's are aligned in parallel with their  
 17 centers at  $I_z = 0$ , then, from Eq.[34], the center of each FPA will be offset from its proper  
 18 position by  $\frac{1}{2}B^2\cot(\text{oga})$ . At the edges of the focal plane ( $B = \pm 0.0070$ ), and at the  
 19 south Pole, the worst-case uncorrected cross-scan offset will be:

$$\begin{aligned} &\text{Worst-case cross-scan} \\ &\text{offset (parallel FPA's)} = \frac{1}{2}(0.0070)^2(1.165) \\ &= 29 \text{ micro-radians} \end{aligned} \quad [39]$$

20 This offset is equivalent to two re-sampled visible pixels (each pixel subtends 14  
 21 micro-radians =  $\frac{1}{2}$  km at nadir) so it is undesirably large. This offset is reduced by  
 22 optimizing Eq. [34] at  $\text{oga} = 45^\circ$ . The worst case cross-scan offset for radial FPA's also

occurs at the edges of the focal plane, for which  $B = \pm 0.0070$ , and at South Pole, where  $B \cot(40.65) = 1.165$ . In this case, the cross-scan offset from the center of the focal plane, at  $B=0$ , to the edges at  $B=\pm 0.007$  is:

$$\begin{aligned} \text{Worst case cross-scan} \\ \text{offset (radial FPA's)} &= \frac{1}{2} (0.0070)^2 (1.165 - 1) \\ &= 4 \text{ micro-radians} \end{aligned} \quad [40]$$

This is 29% of the visible pixels in the imagery, so some re-sampling would be desired to improve worst-case performance for the visible array. It should be noted, however, that the imager's visible range FPA's need not fill the entire along scan dimension of the focal plane, from  $B = -0.007$  to  $B = +0.007$ . The IFOV's at longer wavelengths tend to be larger. Therefore, the worst-case results are probably unduly pessimistic.

If the FPA's are parallel, there is also a time shear in those pixels for which both A and B are non-zero. In other words, if all elements of an array are sampled simultaneously, then the pixels are offset in the along-scan direction. This shear can be derived from the second term of  $I_x(A, B, -B/(2S_o), oga)$  that equals  $-AB \cot(oga)$ . The worst-case error occurs at the corners of the array, where both A and B have values of  $\pm 0.0070$ , and at the south Pole, where  $\cot(oga) = 1.165$ . In this case, the along-scan shear is:

$$\begin{aligned} \text{Along-scan shear (parallel FPA's)} &= \pm (0.0070)^2 (1.165) \\ &= \pm 57 \text{ micro-radians} \end{aligned} \quad [41]$$

This offset is equivalent to about four re-sampled visible pixels. If the FPA's are all parallel to the z-axis in the focal plane, then those FPA's for which B is non-zero will require staggered sampling (like the visible arrays in the GOES Imager) which would be difficult to implement in an array using TDI.



1 Tilting the FPA's in this manner also reduces the worst-case along-scan shear to:

2 
$$\text{Along-scan shear (tilted FPA's)} = \pm(0.0070)^2(1.165-1.00)$$
  
3 
$$= \pm 8 \text{ micro-radians} \quad [42]$$

4 Some correction of this error is required in the re-sampling algorithm to achieve  
5 good co-registration in the Earth imagery in the visible (14 micro-radians) and SWIR (28  
6 micro-radians) channels.

### 7 Preferred Telescope Configuration

8 Referring to Fig. 11, an opto-mechanical system according to another embodiment  
9 of the present invention for implementing a scan pattern across the full Earth disk from  
10 geostationary orbit is illustrated. The primary mirror 1112 and secondary mirror 1114  
11 form a telescope 1110. In contrast to the embodiment of Fig. 6, the optical axis of this  
12 telescope 1110 is tilted 30 degrees from the y-axis in the y-z plane. The along-scan axis of  
13 the FPA 1120 is parallel to the E/W axis of the spacecraft (x-axis). In this system, the oga  
14 is 60 degrees when the LOS is pointed to nadir. The scan mirror 1130 is an optical flat  
15 with an elliptical cross-section that is mounted on a gimbal system 1140 with two  
16 perpendicular axles 1142, 1144. The outer axle 1142 is attached to the body 1102 of the  
17 instrument parallel to its x-axis and perpendicular to the LOS of the telescope. The  
18 projection of the cross-scan axis of the FPA 1120 onto the scan mirror 1130 is always  
19 perpendicular to the outer axle 1142 of the gimbal. The outer gimbal axle 1142 is parallel  
20 to the TDI axis of the FPA 1120. The inner axle 1144 of the gimbal 1140 rotates about the  
21 x-axis of the spacecraft and assumes different orientations in the y/z plane. The scan  
22 mirror 1130 is attached to the inner axle 1144 with the normal vector of its mirrored  
23 surface perpendicular to the inner axle 1144. The major axis of its elliptical cross-section  
24 is parallel to the inner axle 1144.

1 The geometry of the Fig. 11 system is preferred to that of Fig. 6 for several  
2 reasons.

3 First of all, the angle of reflection is smaller. Consequently, the reflection  
4 introduces less spurious polarization. Second, the major axis of the scan mirror ellipse can  
5 be made shorter. As a result, the mass and angular momentum of the scan mirror are  
6 reduced. Additionally, rotation of the image is less. This may be stated equivalently as  
7 there being less rotation of the projection of the FPA pixels onto the scene.

8 Referring to Fig. 12, a bi-directional scan pattern **1210** for mapping the full Earth  
9 disk that is generated by the system of Fig. 11 is illustrated. Comparing the scan pattern of  
10 Fig. 12, with that of Fig. 7, the arcs in Fig. 12 are more nearly straight lines. This yet  
11 another way of characterizing that image rotation is reduced according to this preferred  
12 embodiment.

### 13 Working Example

14 A geosynchronous imager according to the present invention has been designed to  
15 satisfy two sets of requirements: those for a next-generation operational weather imager  
16 and those for a geostationary scientific instrument. It is to be mounted on a three-axis  
17 stabilized geostationary satellite. It has 18 spectral channels ranging from 0.47 to 13.3  
18  $\mu\text{m}$ . It scans the full Earth disk, from geosynchronous orbit, in less than 15 minutes while  
19 simultaneously scanning smaller frames at higher rates. To achieve the required coverage  
20 rate, the exemplary geosynchronous imager scans with long linear arrays that subtend  $0.8^\circ$ .  
21 It has a plane scan mirror in object space, mounted on a two-axis gimbal system, followed  
22 by a three-mirror, off-axis, all-reflective telescope. Dichroic beam splitters allow the  
23 geosynchronous imager to have four co-registered focal planes in four different  
24 wavelength bands, each of which has several linear focal plane arrays (FPA's) that are

1 scanned sequentially over the Earth's surface. Both high angular resolution and high  
2 radiometric resolution are satisfied by time delay & integration (TDI) in several of these  
3 FPA's. The outputs from 20 detector elements in a row are summed (N=20) to achieve  
4 adequate radiometric resolution in one spectral band.

5 Because it utilizes TDI, the geosynchronous imager according to this example  
6 cannot tolerate the image rotation that occurs in the prior art GOES scan geometry. In  
7 order to scan with a single plane mirror on a two axis-gimbal, it was necessary to use a  
8 heretofore unknown scanning geometry to maintain the direction of scan parallel to the  
9 TDI axis of the array in spite of image rotation. Scanning geometry according to the  
10 above description overcomes this problem.

11 The geosynchronous imager scans the Earth in a series of east/west arcs whose  
12 curvature is matched to the image rotation in a way that maintains the opto-mechanical  
13 scan velocity along the TDI axis of the focal plane. Each arc is scanned by rotating the  
14 scan mirror at a constant angular velocity about its inner gimbal axis while the outer axis  
15 remains fixed. The scan mirror is rotated about its outer axis during the turn-around time  
16 between scan arcs. The magnitude of the angle of reflection varies by less than  $0.5^\circ$  over a  
17 single arc of the geosynchronous imager's scan pattern. Space measurements taken at the  
18 ends of each arc are used to subtract the instrument's background from the raw Earth data,  
19 so the net signals have only a small bias due to the angular dependence of the scan mirror's  
20 emissivity.

## 21 **Techniques to Mitigate Mis-Registration & Velocity Mis-Match Problems**

22 Additional aspects of the invention include a number of techniques for mitigating  
23 mis-registration and velocity mis-match problems. These aspects are described as follows.

One technique is to use several sub-arrays in each spectral channel to scan the cross-scan dimension of the FPA. Progressively higher electronic TDI scan rates (or larger along-scan detector pitch) are used going from the inner radius of the scanning arc to the outer radius. The TDI scan rate for each sub-array is constant, but is different from that of the other sub-arrays. The effectiveness of this technique is illustrated by considering the example of an imager having a +/- 1.0 deg cross-track FOV with  $N(TDI) = 32$ . The fractional error =  $\text{Acot}(\text{oga}+45)$ , and the worst case occurs at south Pole, where  $\text{oga} = -4.35$  deg.

- For a single TDI rate  $A = 0.0087$  at  $\text{ogi} = 0$  (45 deg), scan rate =  $0.707 * \text{TDI rate}$

$$\text{Fractional error} = (0.0175)(1.165) = 0.0203$$

$$\text{Blur} = 32 * 0.0203 * 0.289 = \mathbf{0.19 \text{ pixel}}$$

- With four sub-arrays, each with a different TDI rate. Set TDI rates @  $\text{ogi} = 0$  (45 deg) Scan rates =  $\text{Sin}(44.25, 44.75, 45.25, 45.75) * (1.414 * \text{nominal TDI rate})$

Considering worst case at the South Pole:

$$\begin{aligned} \text{Mis-match} &= \{S(44.25)S(40.65)\} / \{S(39.9)S(45)\} \\ &= 1.002 \text{ (ratio at center of sub array)} \end{aligned}$$

$$\begin{aligned} \text{Mis-match} &= (0.0175/4)(1.165) \\ &= 0.005 \text{ (from center to edge of sub array)} \end{aligned}$$

$$\text{Worst mis match} = 0.007$$

$$\text{Blur} = 32 * 0.007 * 0.289 = \mathbf{0.065 \text{ pixel}}$$

This represents a factor of three improvement over single FPA.

Another technique is to stagger the timing of tilted sub-arrays. If it is assumed that multiple sub-arrays are used, then staggering the timing of tilted sub-arrays improves along-scan registration when scanning north or south of the central arc. If all arrays are sampled simultaneously in the central arc, then when sampling a smaller arc (S. Pole) the

1 sub-array on the inner radius of the leading channel (the one that is scanned over the FOV  
2 first) should be sampled before the outer sub-array. In the trailing channel, the outer sub-  
3 array should be sampled before the inner sub-array. At the North Pole, the timing is  
4 opposite: The outer leading array and the inner trailing array should be sampled first.

5 Yet another technique is an alternative to the two described above. It is possible to  
6 deliberately introduce distortion into the optical prescription so that the effective focal  
7 length is slightly shorter at the side of the focal plane that images the outer radius of the  
8 arc than it is at the opposite side (that images the inner radius of the arc). This approach  
9 can equalize (at least to first order) the scan velocities on the focal plane, allowing  
10 constant detector pitch and constant TDI rate to be used over a wider array. In this  
11 implementation, multiple arrays would be aligned parallel to one another and not radially  
12 as illustrated in Fig 13.

### 13 **Summary of Advantages**

14 The scanning method and apparatus according to the present invention, including  
15 gimbal geometry, focal plane layout, and associated scanning technique, has a number of  
16 unique advantages over the prior art.

17 A geosynchronous imager according to the present invention can scan its FOR  
18 over the entire surface of the Earth with a single plane scan mirror on a two-axis gimbal.  
19 When the scan motion is implemented, the projection of the focal plane maintains a fixed  
20 orientation with respect to the scan velocity. This feature allows the electronic scan  
21 velocity of a TDI array to be matched to the opto-mechanical scan velocity of the LOS so  
22 that TDI can be implemented over the entire FOR without image blurring. Although this  
23 ideal co-ordination only occurs at the center of the focal plane (there is some), the worst-  
24 case image blurring that results from mis-match at the edges of the focal plane is a small

1 fraction of a pixel's dimension. No prior art imaging scheme can match this combination  
2 of image quality and radiometric quality, minimization of the angular momentum required  
3 to implement a scan, and simplicity of implementation.

4 Another advantage is that the scanning technique according to the present  
5 invention is easy to implement. The outer gimbal axis is offset between scan lines, but  
6 remains stationary during the data-taking portion of each scanning arc, when the inner  
7 gimbal axis rotates at a constant angular velocity (a different velocity for each arc).

8 A further advantage of the present invention is that the magnitude of the reflection  
9 angle varies by only a fraction of one degree during a single scanning arc. This allows  
10 space references to be taken at the beginning and end of each arc at virtually the same  
11 reflection angle that is applicable to the Earth imagery. This feature suppresses shading  
12 due to the angular dependence of the scan mirror's emissivity.

13 Yet another advantage of the present invention is that the radial focal plane layout  
14 allows multiple FPA's to be located in each focal plane. The orientation of the TDI axes  
15 of the FPA's and the co-registration among the FPA's are ideal at the Equator and are quite  
16 acceptable at the poles.

### 17 **Imaging From a Moving Platform**

18 The foregoing portion of this disclosure analyzes an optical scanning method and  
19 apparatus for producing images under conditions where there is no relative motion  
20 between the imager and the scene, i.e., an Earth-viewing imager on a three-axis stabilized  
21 geostationary satellite. However, the scope of the present invention is not limited to such  
22 a scenario. The geostationary examples set forth herein represent imaging when there is  
23 no significant relative motion between the imager and its observation target. These appear

1 to be the most commercially promising uses of the invention but are not its only useful  
2 application.

3 The following discussion addresses imaging from a moving platform. The specific  
4 example is an Earth-viewing, three-axis stabilized satellite in low-Earth orbit (LEO). The  
5 parameters selected for this analysis are arbitrary. The same approach is generally useful  
6 for observations from any platform moving at a constant velocity and a constant distance  
7 from the scene to be imaged.

8 It is assumed that the spacecraft yaws to maintain a fixed orientation with respect  
9 to the direction of the ground point velocity, which is due to both the spacecraft's orbital  
10 motion and the Earth's rotation. This angle is at most a few degrees away from the orbital  
11 plane. Of course, the spacecraft not only revolves around the Earth, but also pitches once  
12 per orbit, to maintain a fixed orientation with respect to nadir.

13 The relative motion between the imager's platform and its observational target is  
14 used to scan in the along-track direction. The conical scan, nominally perpendicular to the  
15 velocity of the imager's platform, generates a series of arcs. In order to implement a bi-  
16 directional scan (the most efficient approach), it is necessary to scan the oga from front to  
17 rear during the active scan, while simultaneously scanning the iga. The oga is then slewed  
18 rapidly from rear to front during the turn-around period between arcs. If this motion were  
19 not implemented, there would be gaps and areas of overlap in the coverage pattern.  
20 Accordingly, this embodiment does not have the advantage of maintaining a fixed oga  
21 during the active scan, as is achievable in the above-described geostationary embodiments.  
22 All other advantages of the present invention are applicable to this embodiment, however.  
23 Also, the range of the oga's mechanical scan is small, equal to one-half the width of the  
24 scan line. To enable use of TDI, the focal plane is populated with long linear arrays of

1 photo-detectors. Each linear array has a spectral filter to determine its bandpass. The  
2 different spectral bands are scanned sequentially. It is preferred that the scan velocity of a  
3 point on the ground remain perpendicular to the long axes of the array. This feature  
4 improves the accuracy to which the spectral characteristics of a pixel on the ground, i.e.,  
5 the ratio of radiances among multiple spectral bands, can be determined.

6 The following coordinate system is assumed: The +z-axis points toward nadir; the  
7 motion of the point on the ground is in the -x direction; and the y-axis is perpendicular to  
8 the other two axes, completing a right-handed coordinate system.

9 The telescope's optical axis points forward, 5 degrees toward zenith (an arbitrarily  
10 selected angle) so that its vector is  $[+0.9962, 0, -0.0872]$ . The telescope projects the long  
11 axis (cross-scan) of the FPA in the x-z plane, and the along-scan axis (the TDI-axis, if TDI  
12 is used) parallel to the y-axis. Assuming that the FOV subtends  $\pm 1.5$  deg from the  
13 optical axis in the cross-scan direction, the vectors from the edges of the FOV are:  
14  $[+0.9983, 0, -0.0610]$  and  $[+0.9936, 0, -0.1132]$ .

15 The scan mirror is mounted on a two-axis gimbal. The outer axis of the gimbal is  
16 parallel to the y-axis of the spacecraft. The inner axis is constrained to the x-z plane and  
17 coincides with the long axis of the elliptical cross-section of the plane scan mirror. The  
18 orientation of the telescope and gimbal axes is different in this embodiment than in the  
19 prior one, but the same basic telescope-to-scan mirror geometry is used as in the  
20 geostationary system described above. The scan mirror's normal vector,  $[N_x, N_y, N_z]$  is  
21 defined by:

$$\begin{aligned} N_x &= -\cos(\text{iga})\cos(\text{oga}) \\ N_y &= +\sin(\text{iga}) \\ N_z &= +\cos(\text{iga})\sin(\text{oga}) \end{aligned}$$



1 The oga is defined as the angle between the inner gimbal axis and the -x-axis of the  
2 spacecraft. The iga is defined as the angle between the normal vector and the x-z plane.

3 In the home position, iga = 0 degrees and oga = 45 degrees (the oga value having  
4 been selected arbitrarily). The normal vector at the home position is [-0.707, 0, +0.707].  
5 In this position, the telescope's optical axis is reflected at an angle 5 degrees aft of the  
6 nadir. [-0.0872, 0, +0.9962]. The edges of the long axis of the array (cross-scan) are also  
7 reflected in the x-z plane, to [-0.0610, 0, +0.9983] and [-0.1132, 0, +0.9936]

8 The system scans in arcs, by rotating the iga in alternating directions, from -10 to  
9 +10 degrees, then from +10 to -10 degrees, etc. The iga is scanned at an angular rate of 4  
10 deg/sec, taking 5 seconds to complete each arc. Approximately 0.6 seconds is allowed for  
11 the mirror to reverse its direction of rotation about the iga axis. During the active portion  
12 of the scan, the gimbal rotates about the oga axis (y-axis) to compensate for the motion of  
13 the spacecraft. At nadir, a point on the Earth's surface 5 degrees aft of the spacecraft has  
14 an apparent motion of 0.54 deg/sec in the -x direction. Therefore, the oga is scanned at a  
15 rate of -0.27 deg/sec. This oga rate is variable as a function of the iga angle and latitude of  
16 the Earth to match the ground track rate.)

17 Rotation about the oga produces an optical angular motion that is twice the oga  
18 rotation. Therefore, the oga is only required to rotate +/- 0.675 degrees about its nominal  
19 position. Flex pivots can be used instead of an axle. Flex pivots eliminate the mechanical  
20 problems associated with a bearing and conduct heat more efficiently than bearings. It is  
21 also easy to use a cable to transfer power, data, etc. across the outer gimbal axis.

22 The average scan velocity of the line-of-sight from each pixel is about 6.15 deg/sec  
23 = 107 mrad/sec. The dwell time for a Vis/NIR pixel that subtends 14.2 micro-radians is  
24 132 micro-seconds. A TDI of 4-8 provides sufficient integration time on each pixel. The

1 array is configured to perform the TDI integration in both directions, alternating on  
2 consecutive scan lines.

3 The foregoing imaging system is appropriate for imaging to meet the needs of a  
4 next generation visible/near IR remote sensing instrument, i.e., the next generation to the  
5 Thematic Mapper (TM) system, for a spacecraft operating according to the following  
6 parameters:

TO430  
8 Altitude: 700 km +/- 100 km (705 km; same as TM)  
9 Inclination: (98.2 deg; same as TM)  
10 Cross-track swath width: 185-400 km (370 km; twice TM)  
11 Cross-track swath angle = +/- 14.65 deg  
12 Velocity of nadir ground track on Earth's surface = 6.75-6.83 km/sec  
13 Angular velocity of nadir ground track = 9.57-9.69 mrad/sec = 0.548-0.555 deg/sec  
14 Angular velocity of ground track 15 deg fore/aft = 8.82-8.93 mrad/sec  
15 Angular velocity of ground track 15 deg cross-orbit = 9.21-9.32 mrad/sec  
16 Yaw angle between orbital plane and ground track velocity @ equator = 3.8 deg

17 According to these exemplary parameters two satellites are capable of providing  
18 4-day coverage. Spectral band coverage is 4 visible/NIR (silicon photo-detectors) and up  
19 to 3 SWIR. The pixel size is in the ranges of 10-15 m for VNIR and 20-30 m for SWIR.  
20 The detector pitch is 12 microns for VNIR arrays (typical) and 24 microns for SWIR  
21 arrays (typical). The Angular IFOV @ nadir is 14.2 micro-radians and the dwell  
22 time/pixel @ nadir is 1.47 msec. The minimum aperture diameter (assuming that the Airy  
23 disk diam = IFOV @ 0.85  $\mu$ m) is 146 mm (or 5.75 in.). The EFL = 846 mm = 33.3 in.  
24 The aperture is f/# = 5.8. The minimum effective VNIR integration time per pixel (for  
25 desirable SNR) is typically 0.5-1.0 msec, depending on bandwidth. The width of the  
26 image swath is about 37,000 VNIR pixels.

27 According to the foregoing example parameters, the detector element in the center  
28 of the array has a reflection angle that varies from 41.7 deg at the beginning of an arc to 40  
29 deg at the center of the arc to 40.4 degrees at the end of the arc. This is a substantial

1 improvement over an implementation according to the principles of the prior art Thematic  
2 Mapper, wherein the angle of reflection from the scan mirror that is seen by a single  
3 detector element varies by 15 degrees as the mirror scans from one side of the 30 degree-  
4 wide swath to the other. Thus, the angle of reflection from the scan mirror that is seen by  
5 a single detector element varies much less in a conical scan according to the present  
6 invention. This feature is very advantageous in the thermal IR wavelengths where  
7 emissivity of the scan mirror surface as a function of angle is significant.

#### 8 **Alternate Embodiments**

9 The present invention is not limited to the bi-directional scanning. The bi-  
10 directional scan is believed to be the most efficient for most applications. However, the  
11 present invention may be practiced using single-directional scanning for those applications  
12 where it is necessary or appropriate. One way to implement single-directional scanning is  
13 to use the single mirror structure as described above and perform a rapid slew to fly back  
14 between scan arcs. Another option for implementing single-directional scanning is to use  
15 a double-sided mirror and to fly forward between arcs by speeding up rotation of the  
16 mirror during the inactive portion of the scan.

17 The embodiments described above are all in the space context, either geostationary  
18 or low Earth orbit (LEO). However, there is no reason that the structural combinations  
19 and techniques described here should be limited to spaceborne applications. For example,  
20 the imager and imaging techniques according to the present invention may be  
21 implemented for machine vision, in medical imaging, for purposes of pollution detection,  
22 and in the field of intrusion detection (i.e., security).

23 The present invention has been described in terms of preferred embodiments,  
24 however, it will be appreciated that various modifications and improvements may be made

